Exoplanet Discovery Methods

(1) Direct imaging

(2) Astrometry → position wobbles

(3) Radial velocity → velocity wobbles

(4) Transits → “winks”

Today: How to find Earths

(5) Space Transits (Hot Earths)

(6) Gravitational microlensing (Cool Earths)

Later:

(7) Pulsar timing

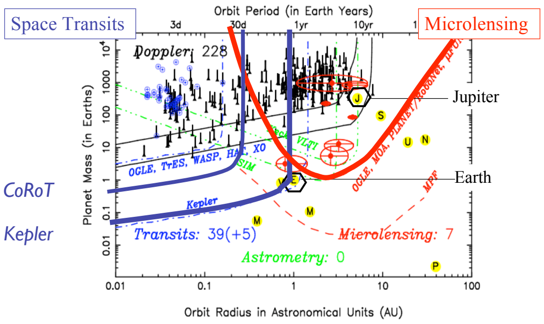
How to find Earths ?

- ***Hot Earths*: Transits from Space**
  - 2007-10 … CoRoT -- Launched Dec 2006
  - 2009-15 … Kepler -- Launched Mar 2009
  - 2017 ? … PLATO
- ***Habitable Earths:* Hard to Find**
  - Habitable Zone: T~300K liquid water on rocky planet surface
- ***Cool Earths*: Gravitational Lensing**
  - 2004-15 … OGLE, MOA, μFUN, PLANET RoboNet

+ KMTNet + LCOGT + SUPA-2 Planet Hunter

“Habitable” Earths: common or rare ?

Exoplanets: 39+228+7=274 (Jun 2008)



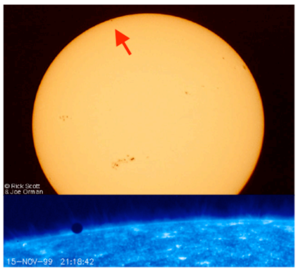
Hot Earths
Transits from Space

Earth transit depth:

$$\frac{\Delta f}{f} \sim 10^{-4} = 0.01\%$$

HST and CoRoT results suggest this is detectable.

Mercury transiting the Sun
15 Nov 1999

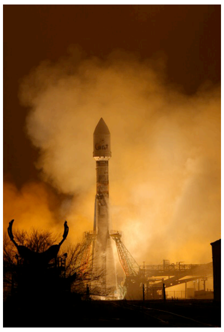


Mercury transits:
2003 May 07
2006 Nov 08

Venus transits:
2004 Jun 08
2012 Jun 06

CoRoT (CNES/ESA)

Launch 27 Dec 2006

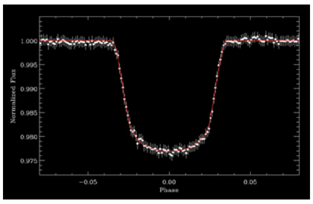


First CoRoT planet:
3 May 2007

6 months/field

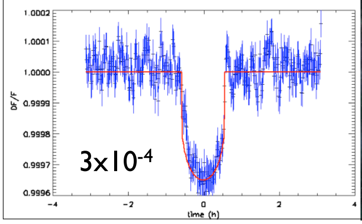
$$\sigma \sim 10^{-3} (t/\mathrm{min})^{1/2}$$

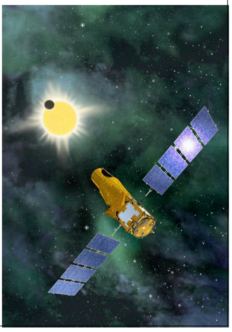
CoRoT-Exo-1b:
P = *1.5* d
m sin(i) = *1.3* m

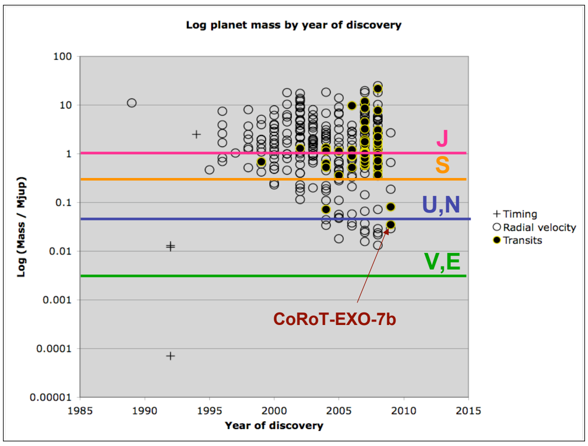


CoRoT-EXO-7b
Smallest Transiting Planet

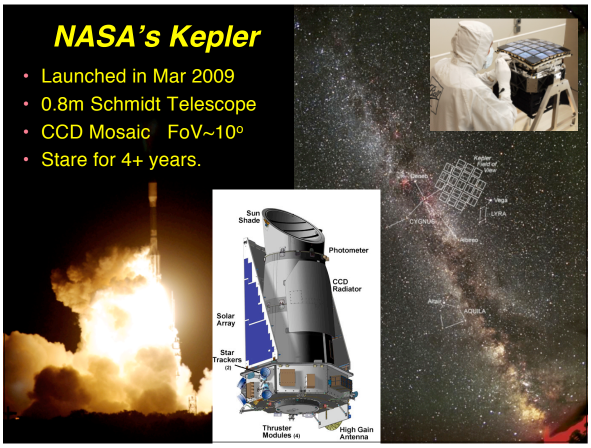
CNES/ESA CoRoT Mission
First transiting Super-Earth
Announced 2009 Feb 3
Planet radius **1.8 Earth radii**
Planet mass 6 - 11 Earth masses





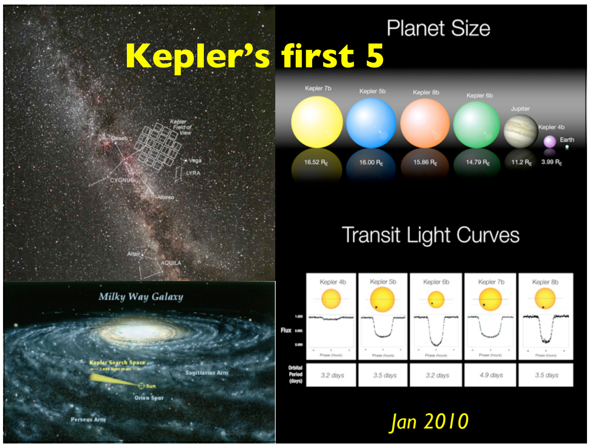


Log planet mass by year of discovery



NASA's Kepler

- Launched in Mar 2009
- 0.8m Schmidt Telescope
- CCD Mosaic FoV~10°
- Stare for 4+ years.

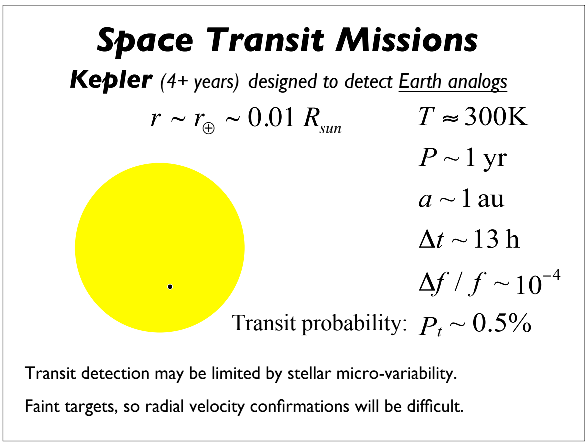


Kepler's first 5

Planet Size

Transit Light Curves

Jan 2010



Space Transit Missions

Kepler *(4+ years) designed to detect Earth analogs*

$r \sim r_\oplus \sim 0.01\, R_{sun}$

$T \approx 300\mathrm{K}$

$P \sim 1\ \mathrm{yr}$

$a \sim 1\ \mathrm{au}$

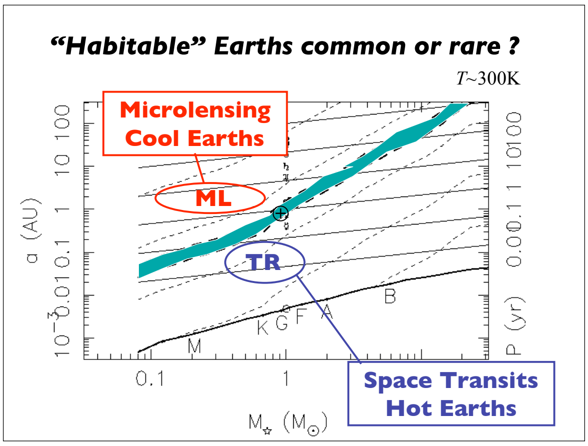
$\Delta t \sim 13\ \mathrm{h}$

$\Delta f / f \sim 10^{-4}$

Transit probability: $P_t \sim 0.5\%$

Transit detection may be limited by stellar micro-variability.

Faint targets, so radial velocity confirmations will be difficult.

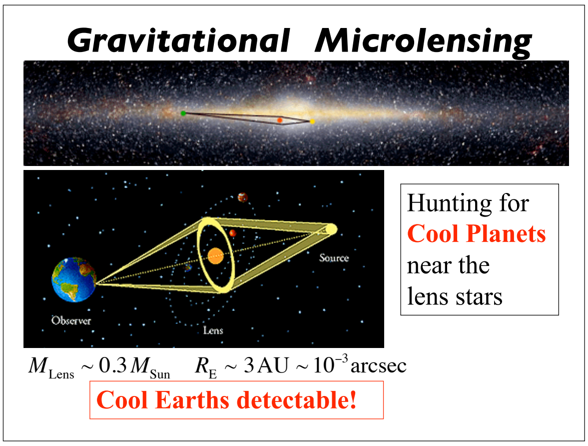


"Habitable" Earths common or rare ?

T~300K

Microlensing Cool Earths

Space Transits Hot Earths



Gravitational Microlensing

Hunting for
Cool Planets
near the
lens stars

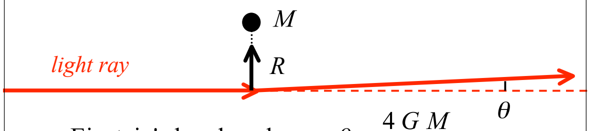
$M_{\mathrm{Lens}} \sim 0.3\, M_{\mathrm{Sun}}$ $R_{\mathrm{E}} \sim 3\,\mathrm{AU} \sim 10^{-3}\,\mathrm{arcsec}$

Cool Earths detectable!

Einstein's General Relativity

Particles (and light) follow shortest available paths (geodessics) through Space-Time.

Mass (energy) causes Space-Time to warp.

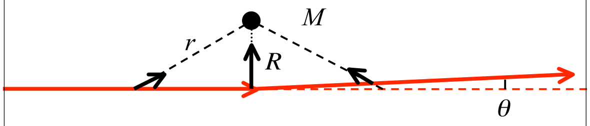


Einstein's bend angle $\theta = \dfrac{4\,G\,M}{R\;c^2}$

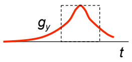
Predicts 1.7 arcsec for Sun-grazing ray.

Verified by Eddington during solar eclipse.

Newtonian Deflection Angle



vertical acceleration $g_y = \left(\dfrac{G\,M}{r^2}\right)\left(\dfrac{R}{r}\right) \le \dfrac{G\,M}{R^2} = g_{\max}$

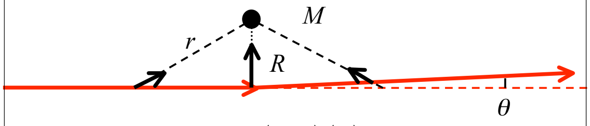


time to pass $\Delta t \approx 2\,R/V_x$

vertical velocity $V_y = \int g_y\,\mathrm{d}t \approx g_{\max}\Delta t \approx \left(\dfrac{G\,M}{R^2}\right)\left(\dfrac{2\,R}{V_x}\right) = \dfrac{2\,G\,M}{RV_x}$

bend angle $\theta \approx \dfrac{V_y}{V_x} \approx \dfrac{2\,G\,M}{RV_x^{\,2}} \Rightarrow \dfrac{2\,G\,M}{Rc^2}$

Newtonian Deflection Angle

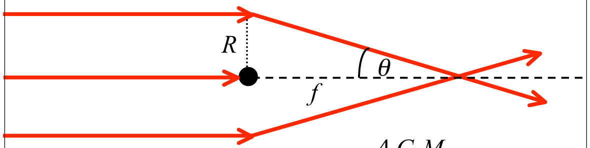


vertical acceleration $g_y = \left(\dfrac{G\,M}{r^2}\right)\left(\dfrac{R}{r}\right)$ $r^2 = R^2 + x^2$

vertical velocity $V_y = \int g_y\,dt \approx \int \dfrac{G\,M\,R}{\left(R^2+x^2\right)^{3/2}}\dfrac{dx}{V_x} = \dfrac{2\,G\,M}{RV_x}$

bend angle $\theta \approx \dfrac{V_y}{V_x} \approx \dfrac{2\,G\,M}{R\,V_x^{\,2}} \Rightarrow \dfrac{2\,G\,M}{R\,c^2}$

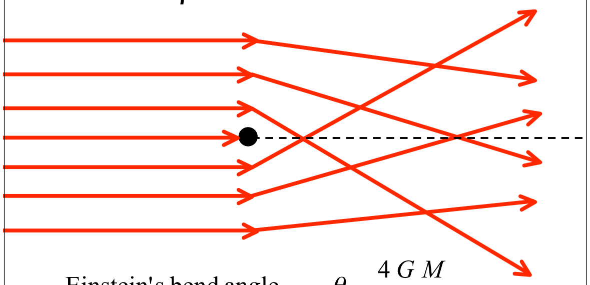
Focal Length of Gravitational Lens



Einstein's bend angle $\theta = \dfrac{4\,G\,M}{R\;c^2}$

Focal length : $f = \dfrac{R}{\theta} = \dfrac{R^2\,c^2}{4\,G\,M}$

Spherical Aberration

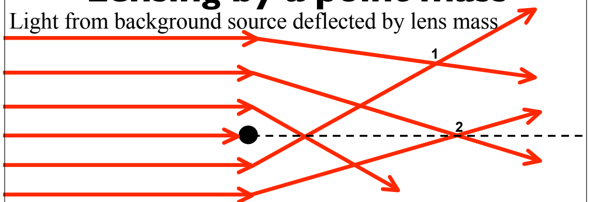


Einstein's bend angle $\theta = \dfrac{4\,G\,M}{R\;c^2}$

Focal length : $f = \dfrac{R}{\theta} = \dfrac{R^2\,c^2}{4\,G\,M}$

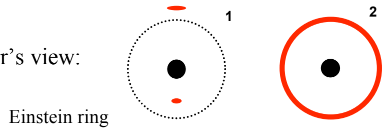
Lensing by a point mass

Light from background source deflected by lens mass

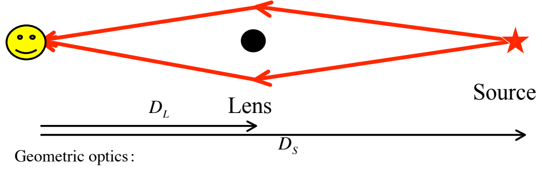


Two distorted/magnified images of background source

Observer's view:

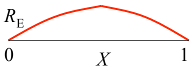


Einstein Ring Radius



Geometric optics:

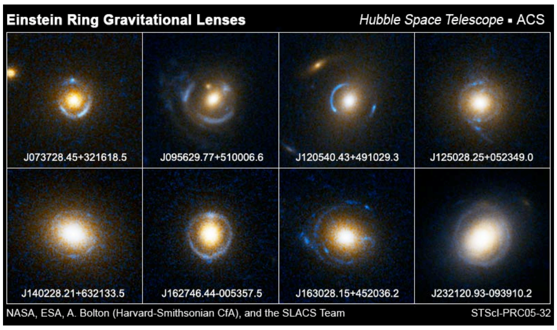
$$\frac{1}{D_S - D_L} + \frac{1}{D_L} = \frac{1}{f} = \frac{4\,G\,M}{c^2\,R^2}$$



Einstein Ring Radius: $X \equiv \frac{D_L}{D_S}$

$$R_E = \left(\frac{4\,G\,M}{c^2}\right)^{1/2} D_S^{\,1/2} \sqrt{X(1-X)} = 8\,\mathrm{AU} \left(\frac{M}{M_{SUN}}\right)^{1/2} \left(\frac{D_S}{8\,\mathrm{kpc}}\right)^{1/2} \sqrt{X(1-X)}$$

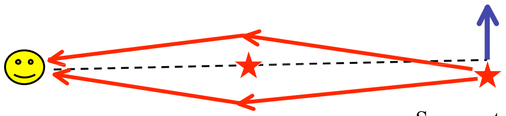
$$\theta_E = \frac{R_E}{D_L} = \sqrt{\frac{4\,G\,M}{c^2}\left(\frac{1}{D_L} - \frac{1}{D_S}\right)} \sim 10^{-3}\,\mathrm{arcsec} \left(\frac{M}{M_{SUN}}\right)^{1/2} \left(\frac{D_S}{8\,\mathrm{kpc}}\right)^{-1/2} \sqrt{\frac{1-X}{X}}$$


Einstein Ring Gravitational Lenses *Hubble Space Telescope* • ACS

NASA, ESA, A. Bolton (Harvard-Smithsonian CfA), and the SLACS Team STScI-PRC05-32

Real images of Einstein Rings. The bright yellow object is the foreground "lens". The blue arcs are images of the background "source". The images of the source form in a ring around the lens.

Lensing of Stars by Stars



V_{rel} ~ 200 km/s

Lens star, possibly
with planets

Source star
near centre of
Milky Way

Bend angle ~ 1 milli-arcsec

~500 cases found every year

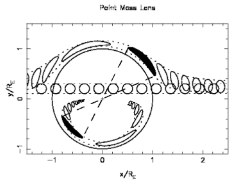
Lensing by a Point Mass

2 images

opposite sides of lens

major image outside ring

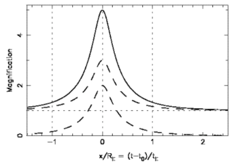
minor image inside ring



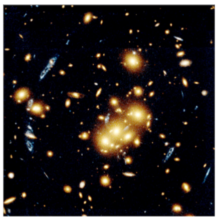
net magnification

(sum of 2 images)

vs time

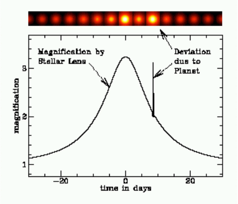


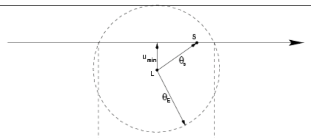
Mass of the Lens determines the type of Gravitational Lensing



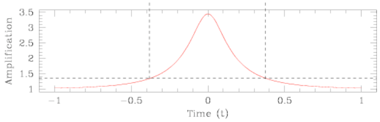
Macrolensing: Very massive lenses, like galaxies or galaxy clusters, produce resolvable images of background objects.

Microlensing: A small lens, like a single star, produces unresolved images. We observe a time-dependent brightening of the source.





Dimensionless source-lens separation: $u \equiv \theta_S / \theta_E$



The total magnification (sum of both images) :

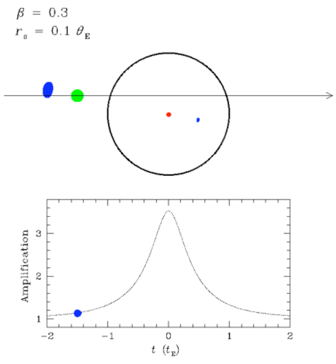
$$A = \frac{u^2 + 2}{u\sqrt{u^2 + 4}} \qquad u(t) = \left(u_{\min}{}^2 + \left(\frac{t - t_0}{t_E}\right)^2 \right)^{1/2}$$

At closest approach, $t = t_0$, $u(t_0) = u_{\min}$

“*Einstein time scale*”, t_E, is the time for the source to move, with respect to the lens, by one Einstein ring radius

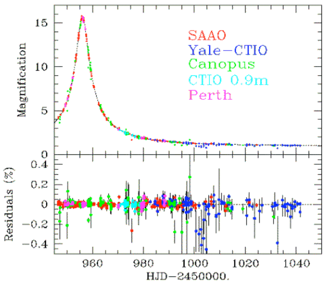
$$t_E \equiv \frac{R_E}{V_{rel}} \propto \sqrt{M}$$

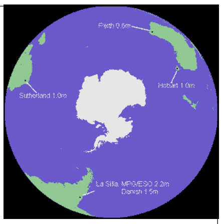
V_{rel} is the relative velocity between the source and lens in the plane of the sky.



Microlens Follow-up Network:

OGLE, MOA, μFUN, PLANET / RoboNet





Global Telescope Network

=>24-hour coverage

1 day => Jupiters

1 hour => Earths

1990s, several groups monitored Galactic Bulge and the Magellanic Cloud starfields to detect lensing by foreground objects (*MACHO, EROS, MOA, OGLE*).
Original motivation was to search for Dark Matter in the form of massive compact halo objects (MACHOs).

Timescales $t_E \sim M^{1/2}$ for Galactic Bulge source stars:

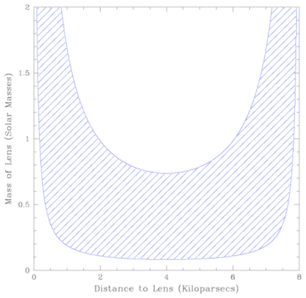
- Solar mass star ~ 1 month
- Jupiter mass planet ~ 1 day
- Earth mass planet ~ 1 hour

These timescales are observationally feasible.

Small bend angles => lensing is a very **rare event**.
Only 1 star in a million is lensed at any given time.

Galactic Bulge surveys (OGLE, MOA) find ~ 600 events/ year.

Degeneracy of Lens Parameters



- Event timescale t_E is a function of lens mass M_L, distance D_L, and relative velocity V_{rel}.
- A continuum of lens parameters can produce the same t_E
- For t_E = 40 d and V_{rel} = 100 - 300 km/s, M_L and D_L can be anywhere in the shaded region.

Notes:
(1) Peak magnification depends on the impact parameter, small impact parameter -> large magnification (A~1/u).

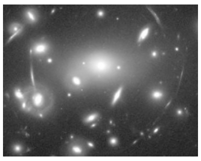
(2) For u = 0, apparently infinite magnification! In reality, finite size of source star limits the peak magnification.

(3) Significant magnification (A>1.3) requires alignment smaller than the Einstein ring radius (u<1).

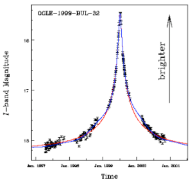
(4) Microlensing is achromatic – all wavelengths affected equally.

(5) Chances of microlensing occurring for a particular star is around 1 in a million – any given star lensed only once.

The Mass Profile of the Lens affects the structure of the source images (macrolensing) or shape of lightcurve (microlensing)

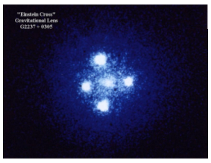


Lens = Galaxy Cluster



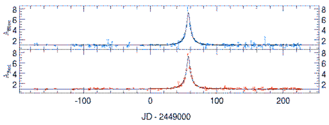
Lens = Black Hole?

Note wobble due to Earth’s annual orbit around the Sun.



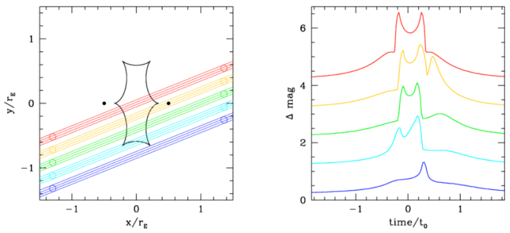
Lens = Single Galaxy

Lens = simple low mass star



Microlensing Anomalies

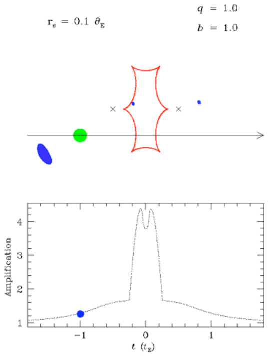
- Deviations from the standard point-lens point-source light curve are referred to as *microlensing anomalies.*
- The most interesting anomaly occurs in the case of a binary lens.
- Lightcurve shape depends on source trajectory relative to binary.



Caustic Crossings

Two new images appear (3->5) when source star enters a caustic curve.

Two images merge and disappear (5->3) when source exits a caustic curve.



How does microlensing find planets?

Light curve for a binary lens is complicated, but a characteristic is the presence of sharp spikes or **caustics**.
With good monitoring, parameters of the binary can be recovered.

Orbiting planet is just a binary with mass ratio $q = m_p/M_* << 1$

- Monitor known lensing events in real-time with dense, high precision photometry from several sites.
- Look for deviations from single star light curve due to planets
- Timescales ~ a day for Jupiter mass planets, ~ hour for Earths
- Most sensitive to planets at a ~ R_E, the Einstein ring radius
- Around 3-5 AU for typical parameters

Aug 2005

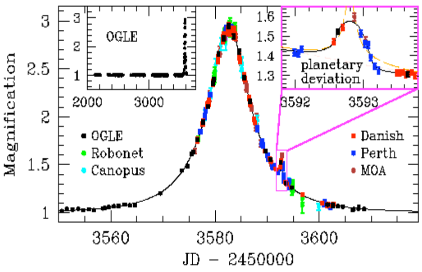
OB-05-390

smallest cool planet

$m \sim 5.5\, m_\oplus$

$a \sim 2.9$ AU

PLANET/
RoboNet
OGLE
MOA



Beaulieu et al. (2006).

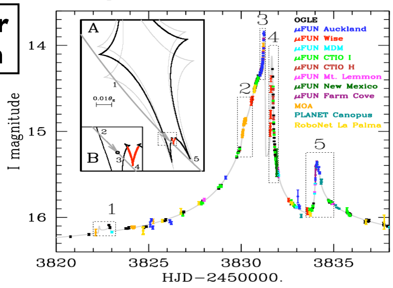
Apr 2006

OB-06-109

2 planets detected !

m_b~jupiter
m_c~saturn

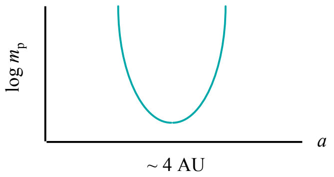
OGLE
microFUN
MOA
PLANET/
RoboNet



Gaudi et al. (2008)

Sensitivity to "Cool" Planets

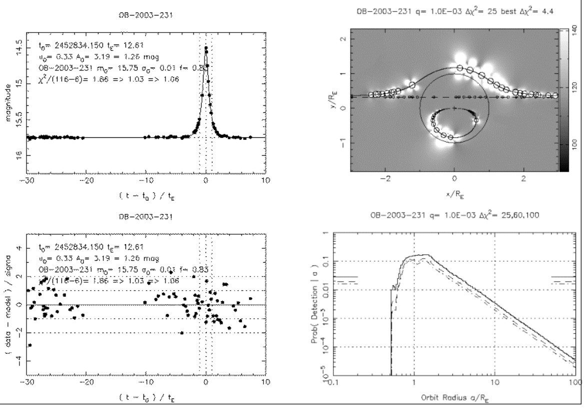
Complementary to other methods:

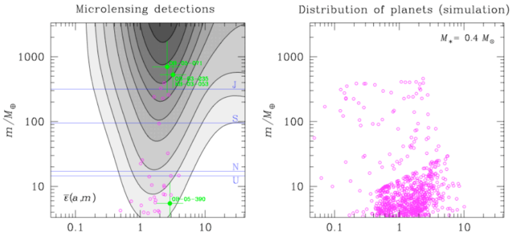


Sensitivity is hard to evaluate: depends on **cadence** of photometric monitoring (high cadence needed for lower masses), **accuracy** of photometry (planets produce weak deviations more often than strong ones)

Very roughly: observe with 1-2% accuracy, once per night, detect Jupiters, if present, with 10% efficiency.
Once per hour, detect Earths, if present, with 1% efficiency.

Planet Detection Zones





More (Smaller) Cool Planets to Come!

Mass function of Cool Planets

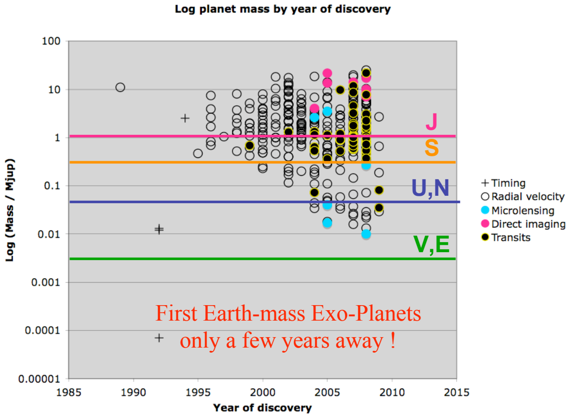
tests Planet Formation / Migration Theory

Microlensing Planets

- 10 planets have been found by microlensing, including a 5.5 Earth-mass object and a multi-planet system.
- Typical lens stars are 0.3 M_{sun} (0.1 - 2 M_{sun})
- Sensitive to “cool” planets at ~ 4 AU, outside the “Snow Line”.
- Determine planet mass, orbit size, star mass, distance.
- **Earth-mass planets can be detected.**

- Monitoring microlensing events takes a lot of time on small telescopes around the world.

- Planet detections are made, but sometimes ambiguous.
- Degeneracy in microlensing light curves -- one lightcurve has several possible solutions for the properties of the source.
- No repeat observations. One time event. Get no more information about the planet-star system.



First Earth-mass Exo-Planets
only a few years away !